

Optimized target allocation algorithm for multi-fibre fed spectrographs MOS@E-ELT

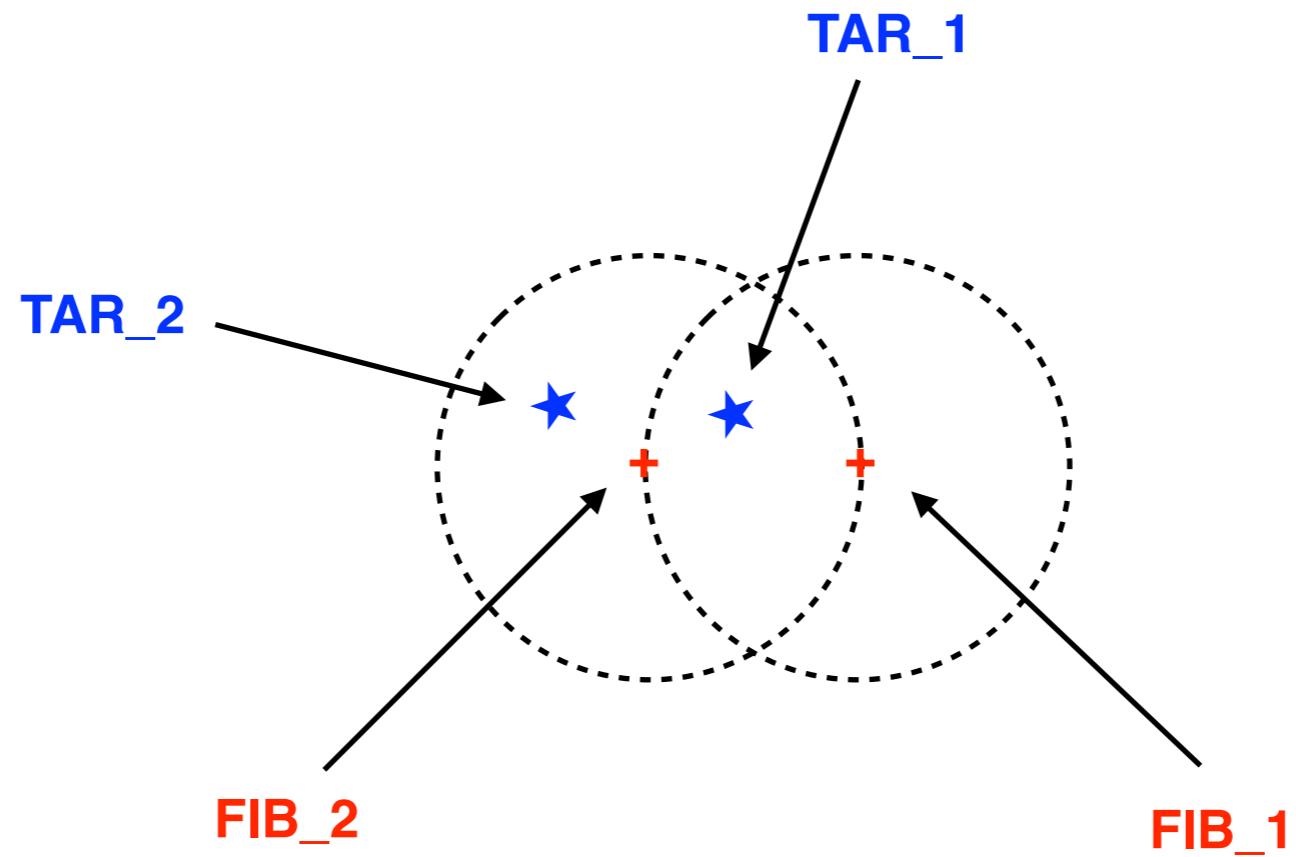
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IASF - Milano

Sexten, July 2015

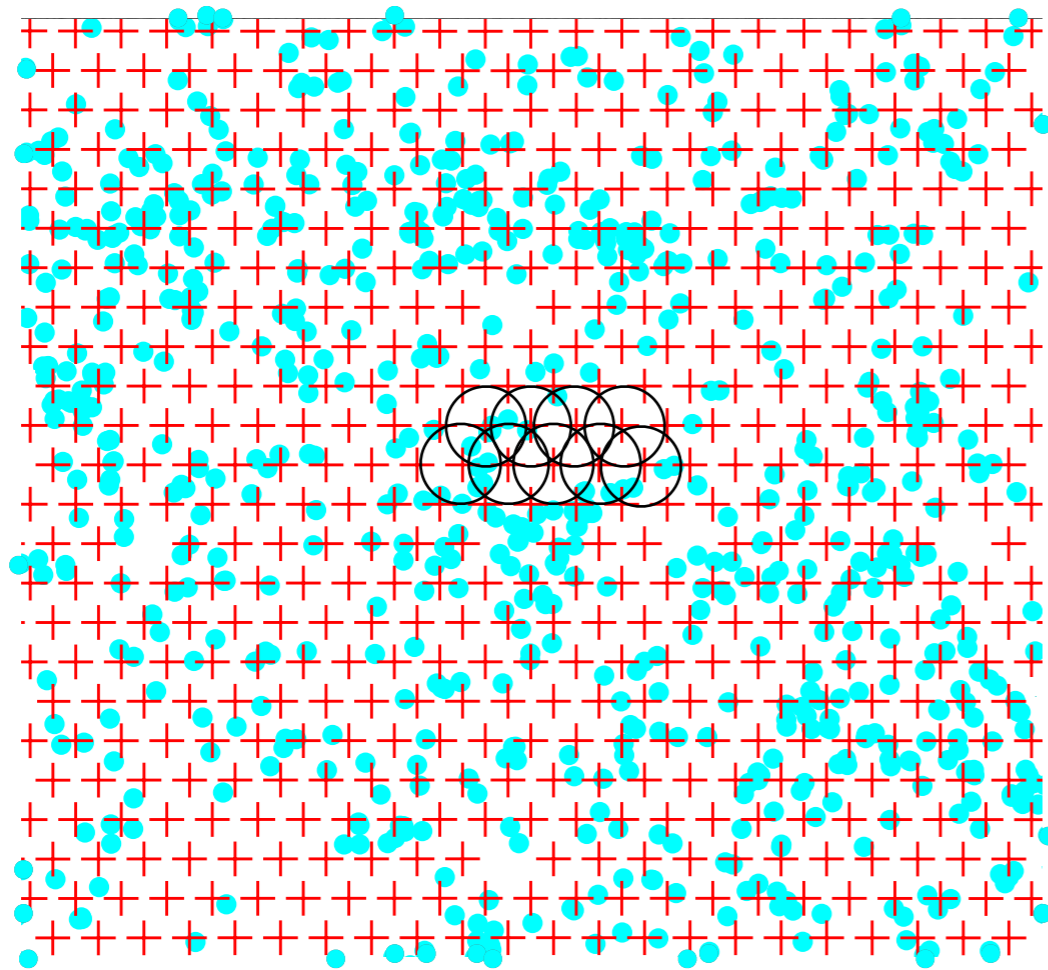


Why the optimization?

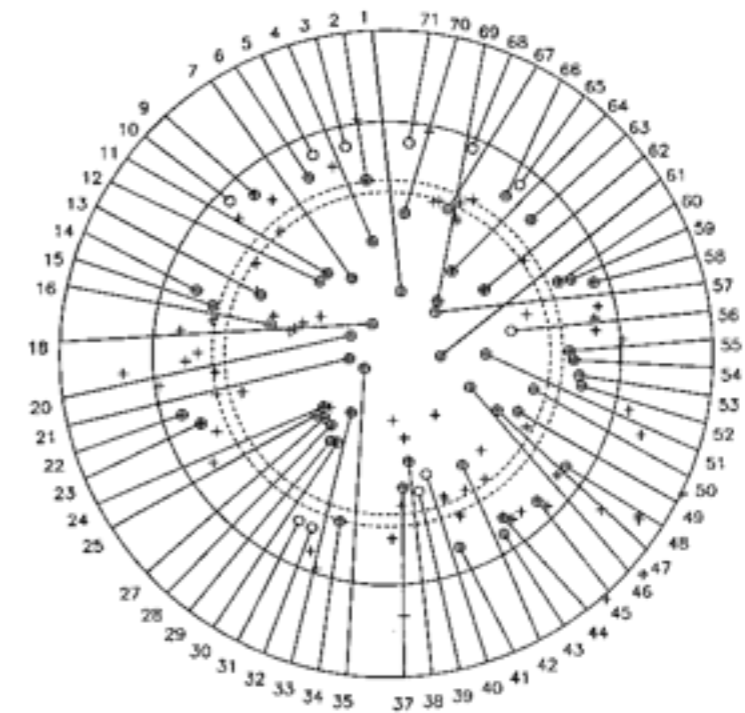


Why the optimization?

Increasing the number of fibres, ~hundreds:
maximizing the allocated targets is a complex problem



AUTOFIB FIELD CONFIGURATION: SEXTANS_APR91_T

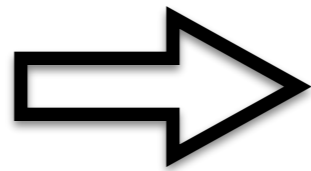
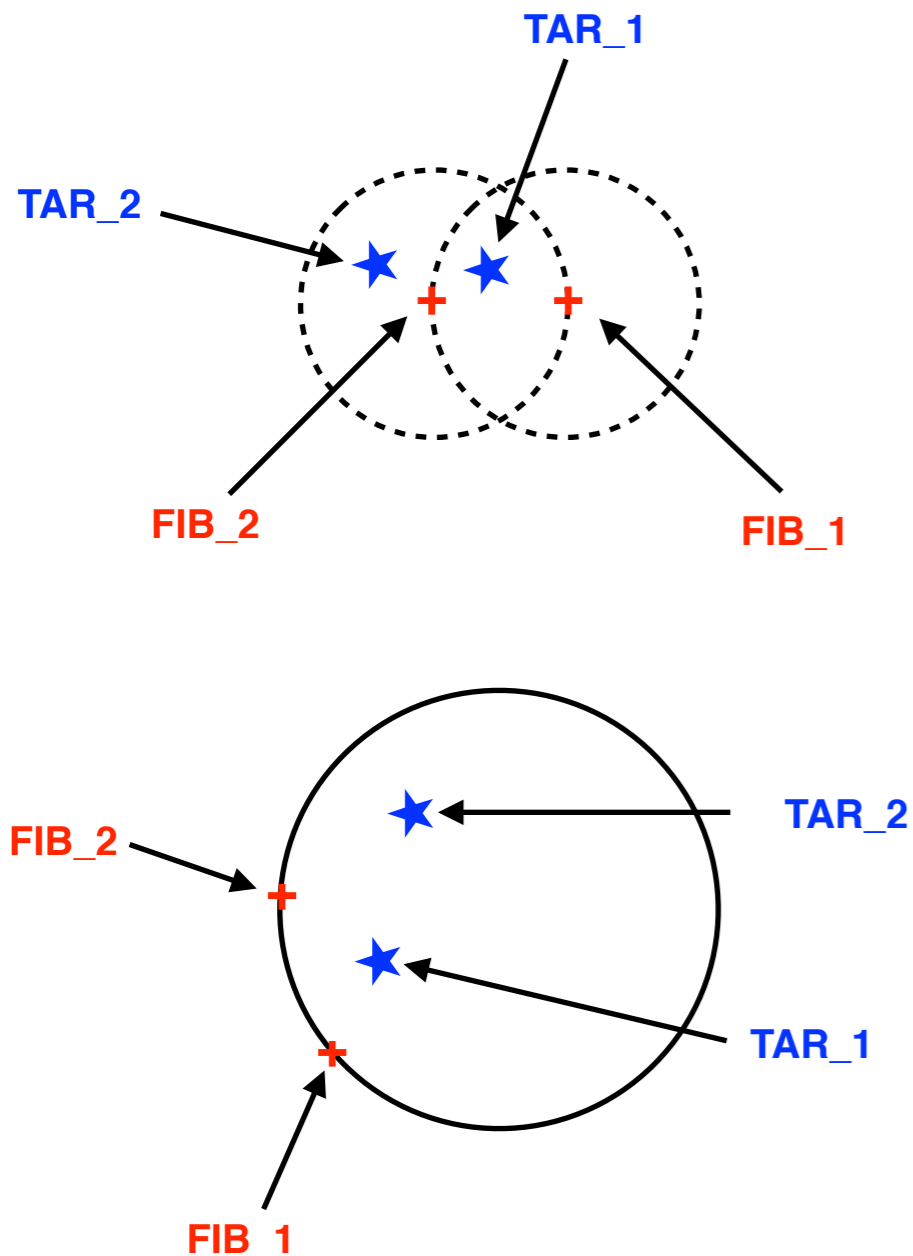


66 Pivots, 130 Objects, TD= 2000, BD= 4500, $\theta_c= 7.5$, $\theta_0= 7.5$, TL= 265000

The optimization **idea**

STORING the INFORMATION:

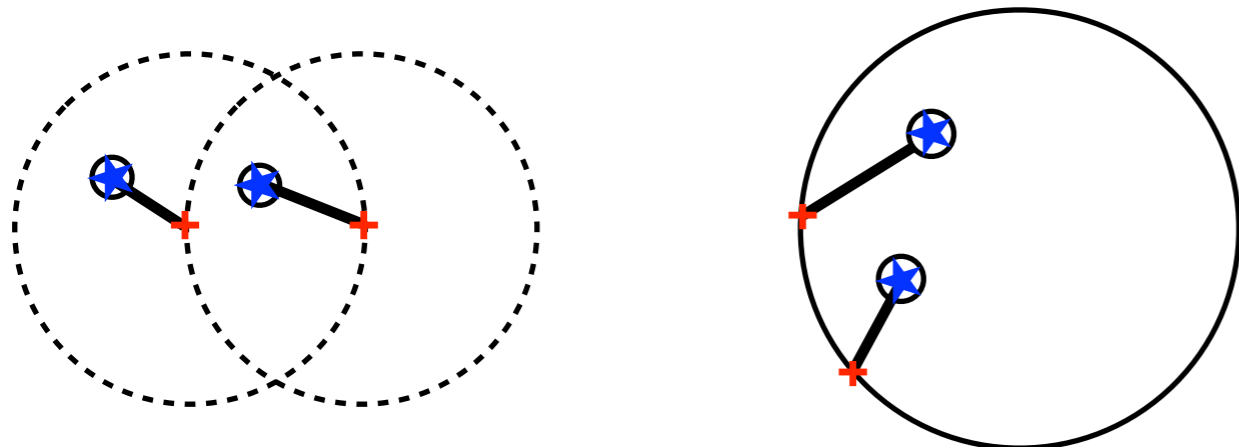
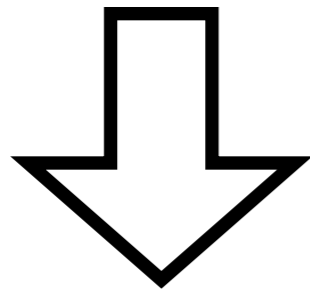
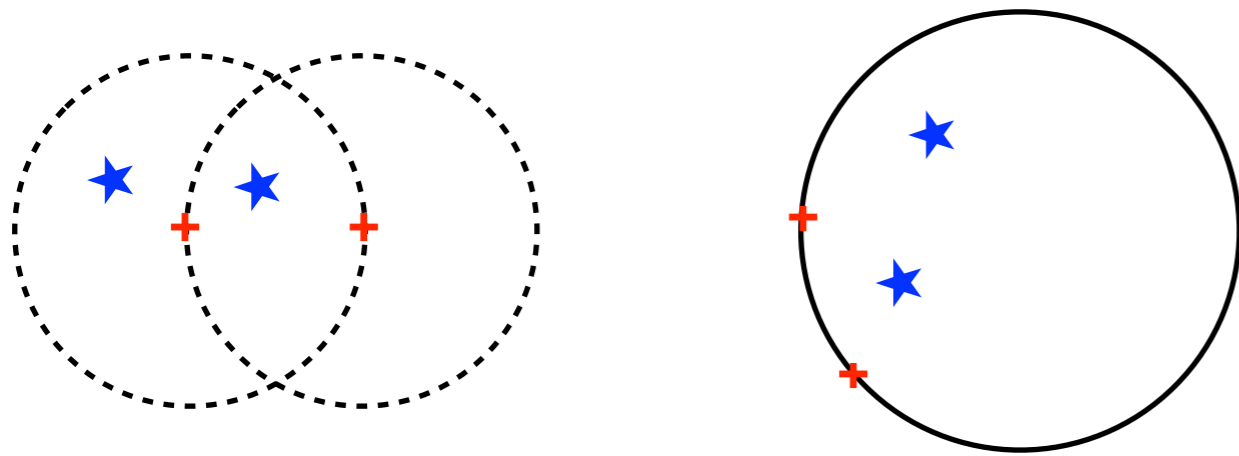
for each fibre, generate a list of targets reachable by that fibre



| FIB_1 | FIB_2 |
|-------|-------|
| TAR_1 | TAR_1 |
| | TAR_2 |

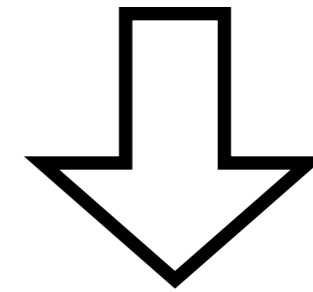
The optimization **idea**

DISCARD targets from the most populated list



BEFORE OPT

| FIB_1 | FIB_2 |
|-------|-------|
| TAR_1 | TAR_1 |
| | TAR_2 |



AFTER OPT

| FIB_1 | FIB_2 |
|-------|-------|
| TAR_1 | TAR_2 |

The optimization **algorithm**

- 1) choose the fibre that can reach more targets,
fibre with more degrees of freedom (REF. Morales et al. 2012)
- 2) select the target reachable by more fibres,
target with more degrees of freedom
- 3) discard the chosen target from the fibre list

| FIB_1 | FIB_2 | ... | FIB_12 | FIB_13 | FIB_14 | FIB_15 | ... |
|-------|-------|-----|-----------|-----------|-----------|-----------|-----|
| | | | TAR_1 (3) | TAR_1 (3) | TAR_1 (3) | TAR_2 (3) | |
| | | | TAR_2 (3) | TAR_2 (3) | TAR_4 (1) | | |
| | | | | TAR_3 (1) | | | |

| ... | FIB_12 | FIB_13 | FIB_14 | FIB_15 | ... |
|-----|-----------|-----------|-----------|-----------|-----|
| | TAR_1 (3) | TAR_1 (3) | TAR_1 (3) | TAR_2 (3) | |
| | TAR_2 (3) | TAR_2 (3) | TAR_4 (1) | | |
| | | TAR_3 (1) | | | |

| ... | FIB_12 | FIB_13 | FIB_14 | FIB_15 | ... |
|-----|-----------|-----------|-----------|-----------|-----|
| | TAR_1 (3) | TAR_1 (3) | TAR_1 (3) | TAR_2 (2) | |
| | TAR_2 (2) | TAR_3 (1) | TAR_4 (1) | | |

| ... | FIB_12 | FIB_13 | FIB_14 | FIB_15 | ... |
|-----|-----------|-----------|-----------|-----------|-----|
| | TAR_1 (2) | TAR_1 (2) | TAR_4 (1) | TAR_2 (2) | |
| | TAR_2 (2) | TAR_3 (1) | | | |

| ... | FIB_12 | FIB_13 | FIB_14 | FIB_15 | ... |
|-----|-----------|-----------|-----------|-----------|-----|
| | TAR_1 (1) | TAR_3 (1) | TAR_4 (1) | TAR_2 (2) | |
| | TAR_2 (2) | | | | |

| ... | FIB_12 | FIB_13 | FIB_14 | FIB_15 | ... |
|-----|-----------|-----------|-----------|-----------|-----|
| | TAR_1 (1) | TAR_3 (1) | TAR_4 (1) | TAR_2 (1) | |

The optimization **output**

OPTIMIZATION OUTPUT

| ... | FIB_12 | FIB_13 | FIB_14 | FIB_15 | ... | FIB_102 | FIB_103 | ... |
|-----|-----------|-----------|-----------|-----------|-----|-----------|-----------|-----|
| | TAR_1 (1) | TAR_3 (1) | TAR_4 (1) | TAR_2 (1) | | TAR_5 (1) | TAR_7 (1) | |
| | | | | | | TAR_6 (1) | TAR_8 (1) | |
| | | | | | | | TAR_9 (1) | |

- After the optimisation each target is assigned to only one fibre
- Selecting targets depending on their priority

Checking for fibre collision

A PRIORI: avoid targets too close because can collide due to the limited size of the fibre

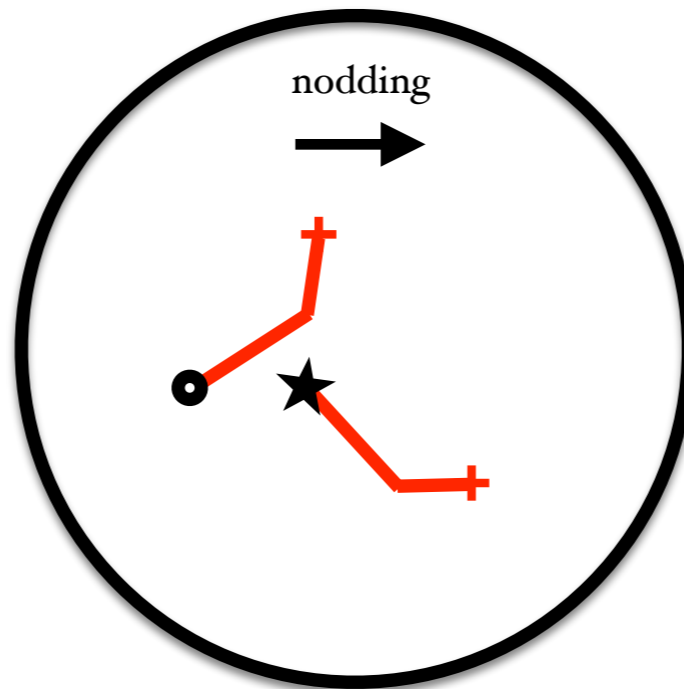


A POSTERIORI: check for fibre collision and, when a conflict occurs, search for other targets



Best sky subtraction

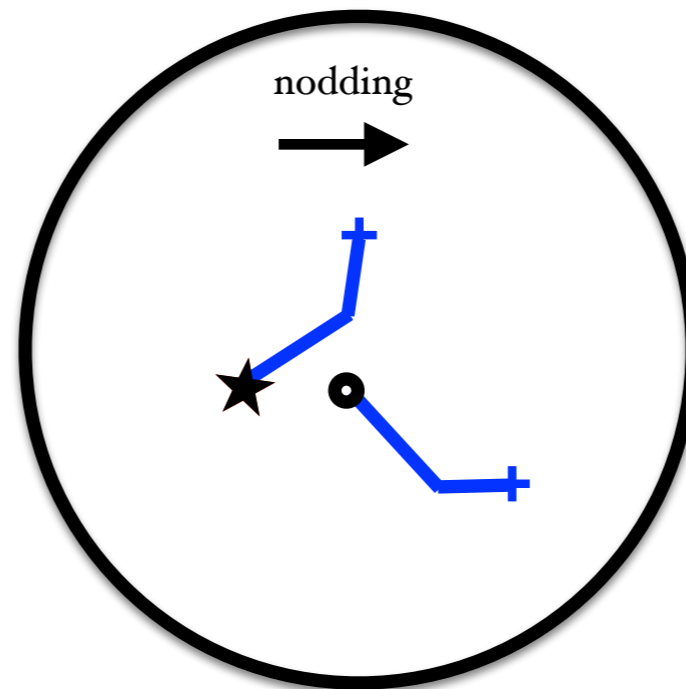
- Infrared observations with high emission of the sky
- Half of the fibres allocated to targets ($\sim 50\%$).
Half ($\sim 50\%$) is left to observe sky positions
- ABBA nodding of the telescope:
target observed during all the observation



Best sky subtraction

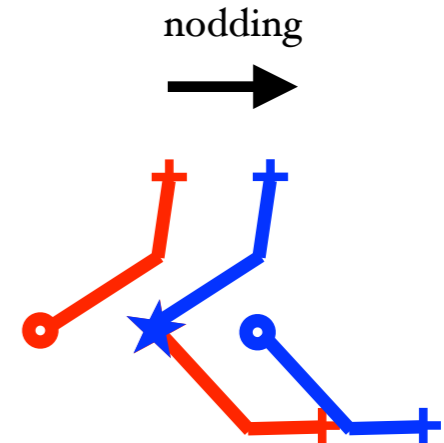
Many objects are lost because:

- targets too close
- the sky position before/after the nodding ends up on a background source



The optimization algorithm

- 1) fix the nodding amount and direction
- 2) store the information
- 3) choose the fibre and randomly select a target
- 4) choose the sky position in the fibre list with the lowest number of targets
- 5) loop over different nodding to reach the best configuration



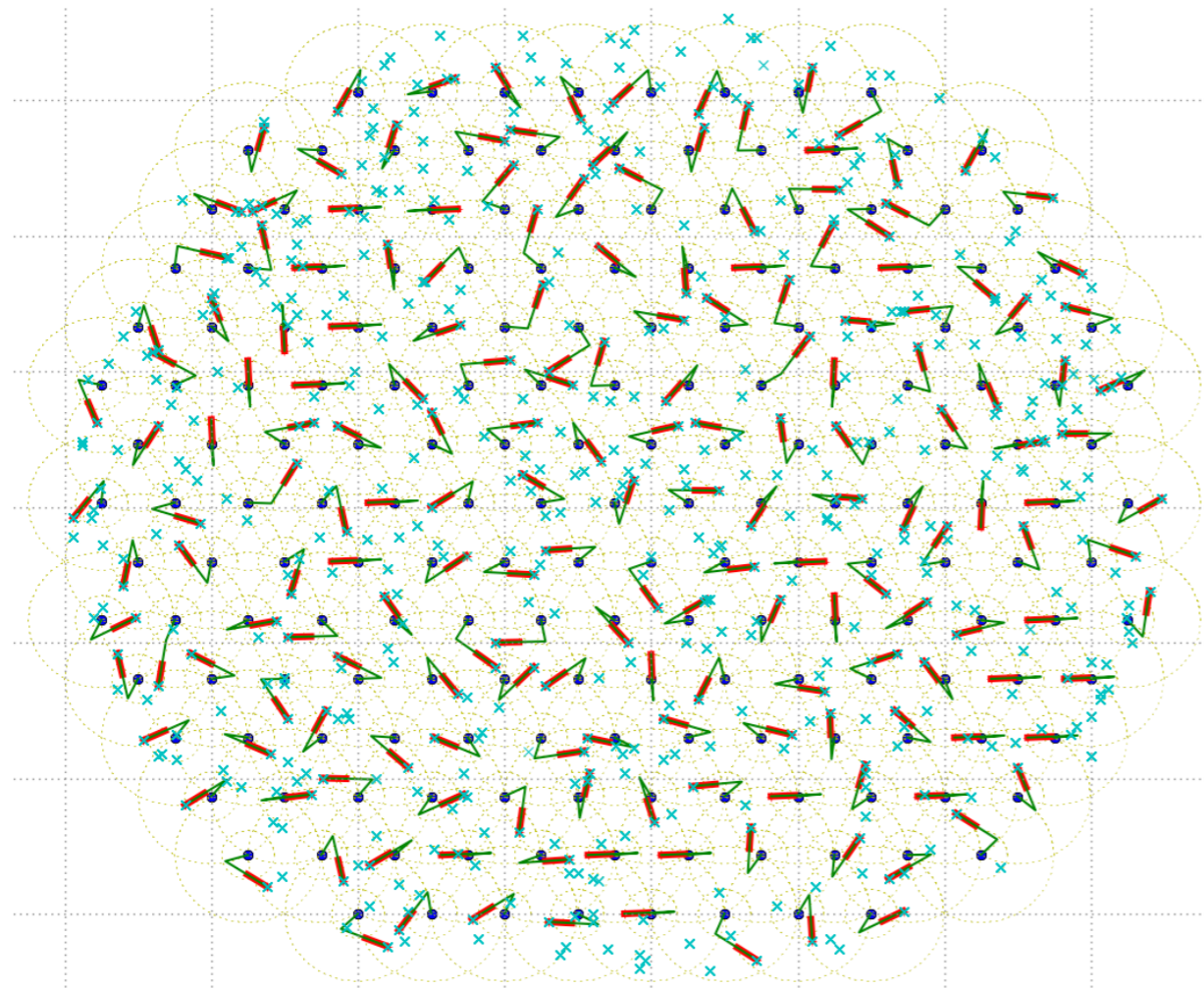
| ... | FIB_12 | FIB_13 | FIB_14 | FIB_15 | ... |
|-----|--------|--------|--------|--------|-----|
| | TAR_1 | TAR_1 | TAR_1 | TAR_2 | |
| | TAR_2 | TAR_2 | TAR_4 | SKY_4 | |
| | SKY_1 | TAR_3 | SKY_4 | | |
| | SKY_2 | SKY_1 | TAR_3 | | |
| | | SKY_2 | SKY_2 | | |
| | | SKY_3 | | | |



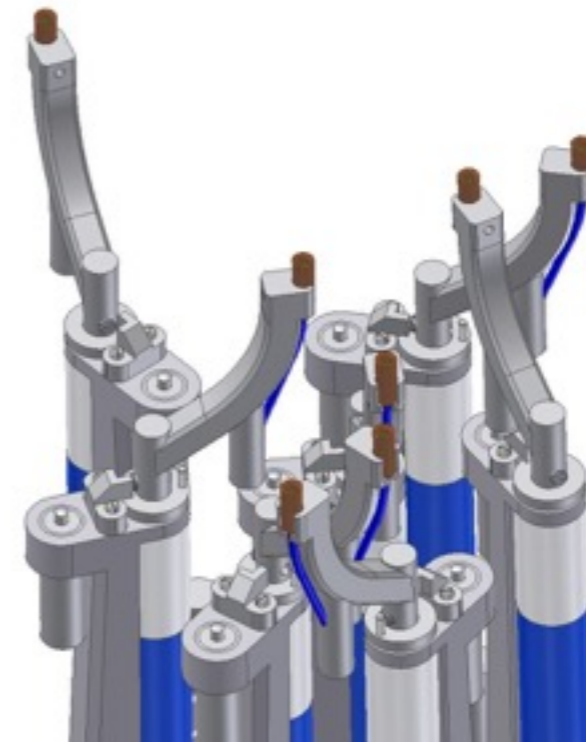
MOONS Spectrograph

3 observing modes:

- Stare mode
- XSwitch mode
- Stare+Nod mode



more than 1000 fibres



Some results ...

INPUT CATALOGUE:

COSMOS photometric redshift catalogue v1.5 (Ilbert et al. 2008)

Bright Sample

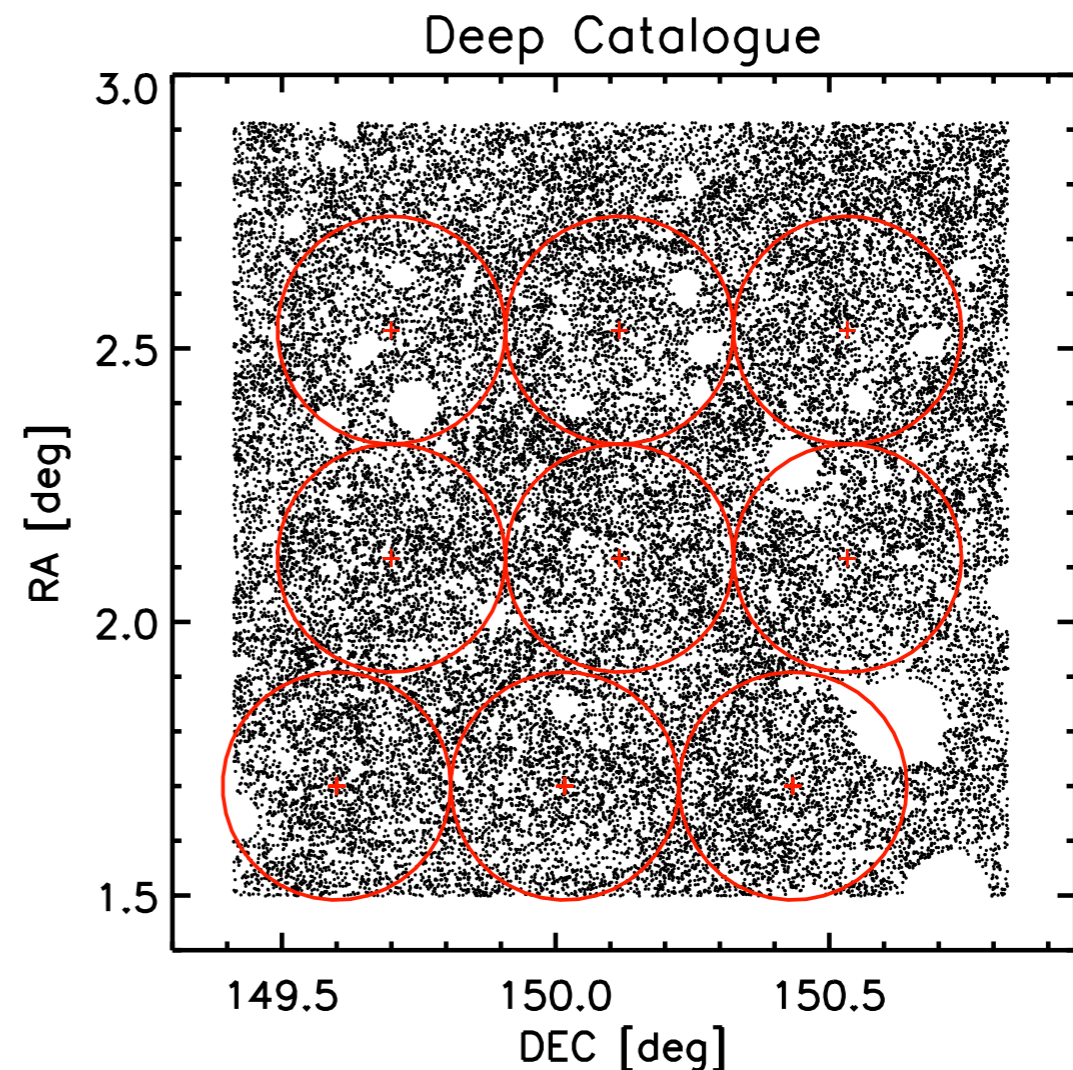
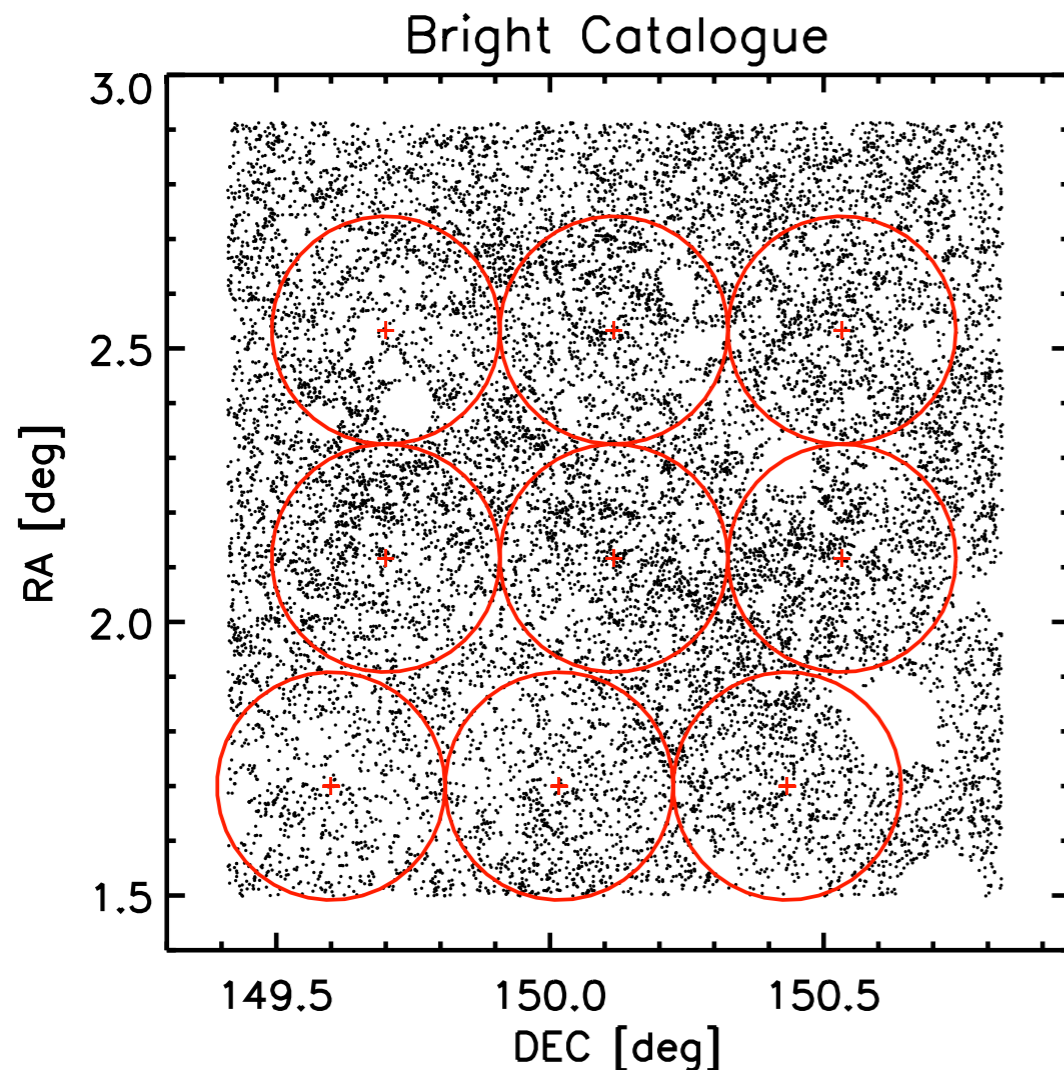
$K = 22.5$ & $0.8 < z < 1.4$

Surface density ~ 1000 per FOV

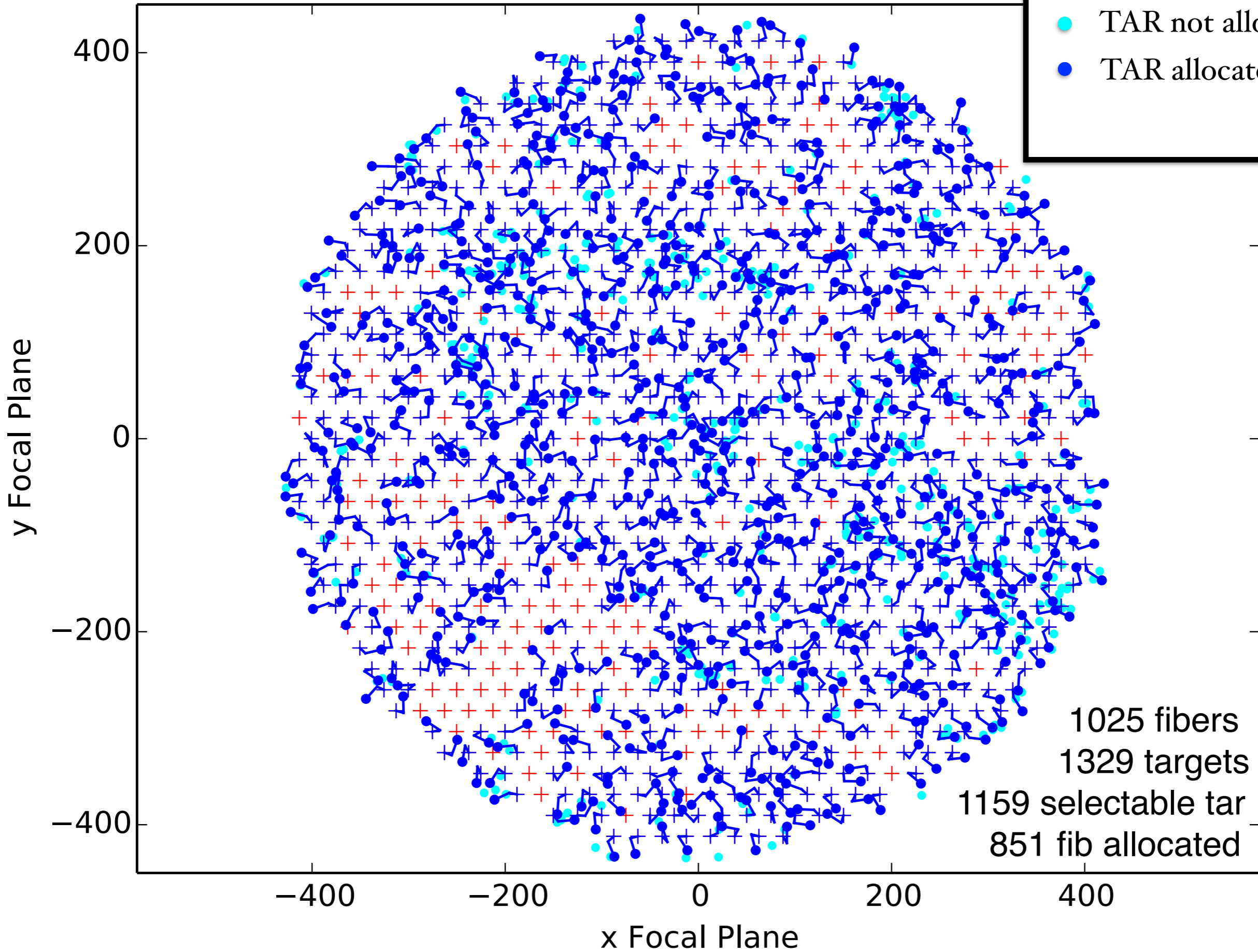
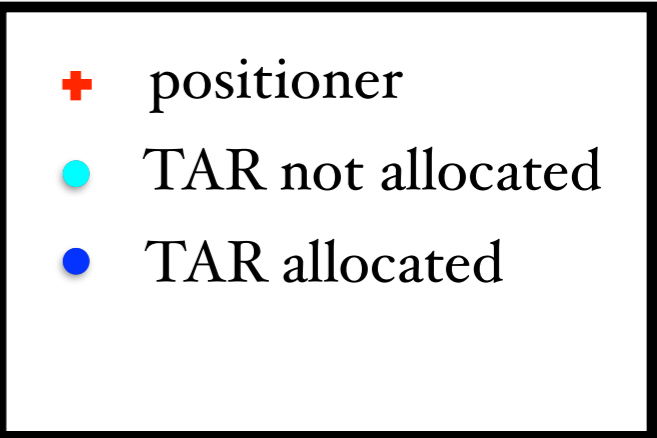
Deep Sample

$K = 23.5$ & $z > 1.4$

Surface density ~ 2500 per FOV

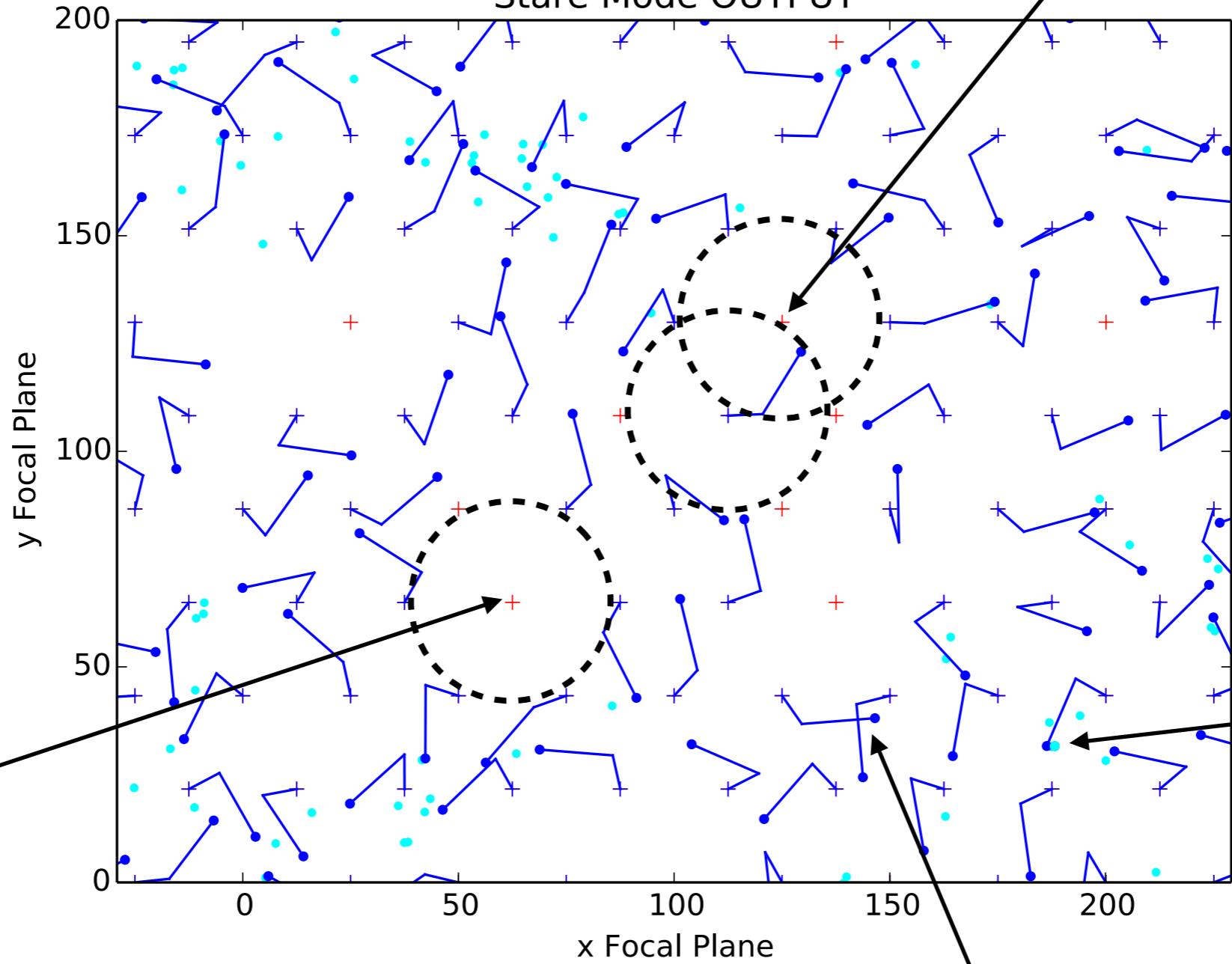


Stare Mode OUTPUT



Inevitable Null
fibre After OPT

Stare Mode OUTPUT



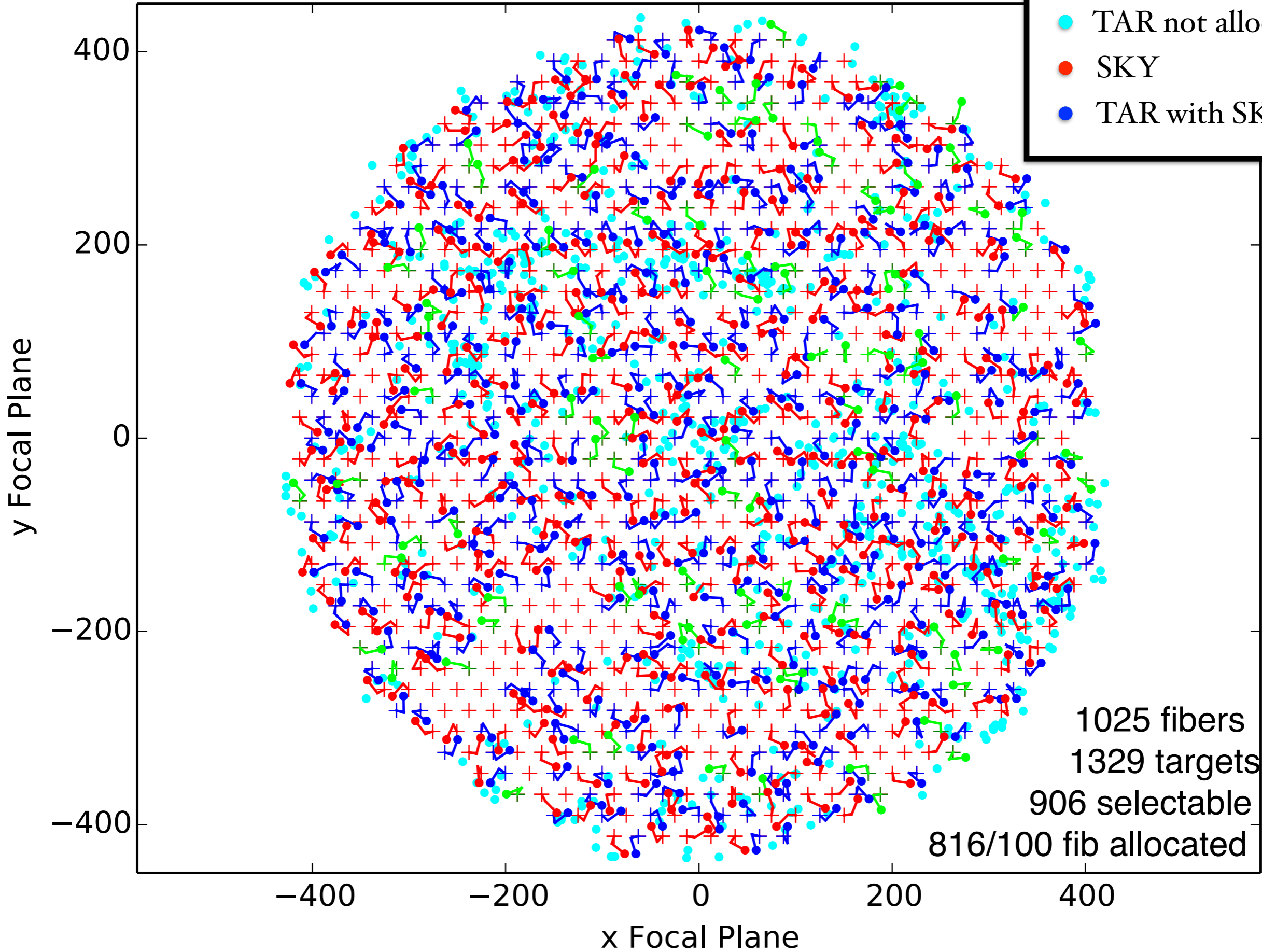
NULL fibre
before OPT

Objects
too close

Collision

XSwitch Mode OUTPUT

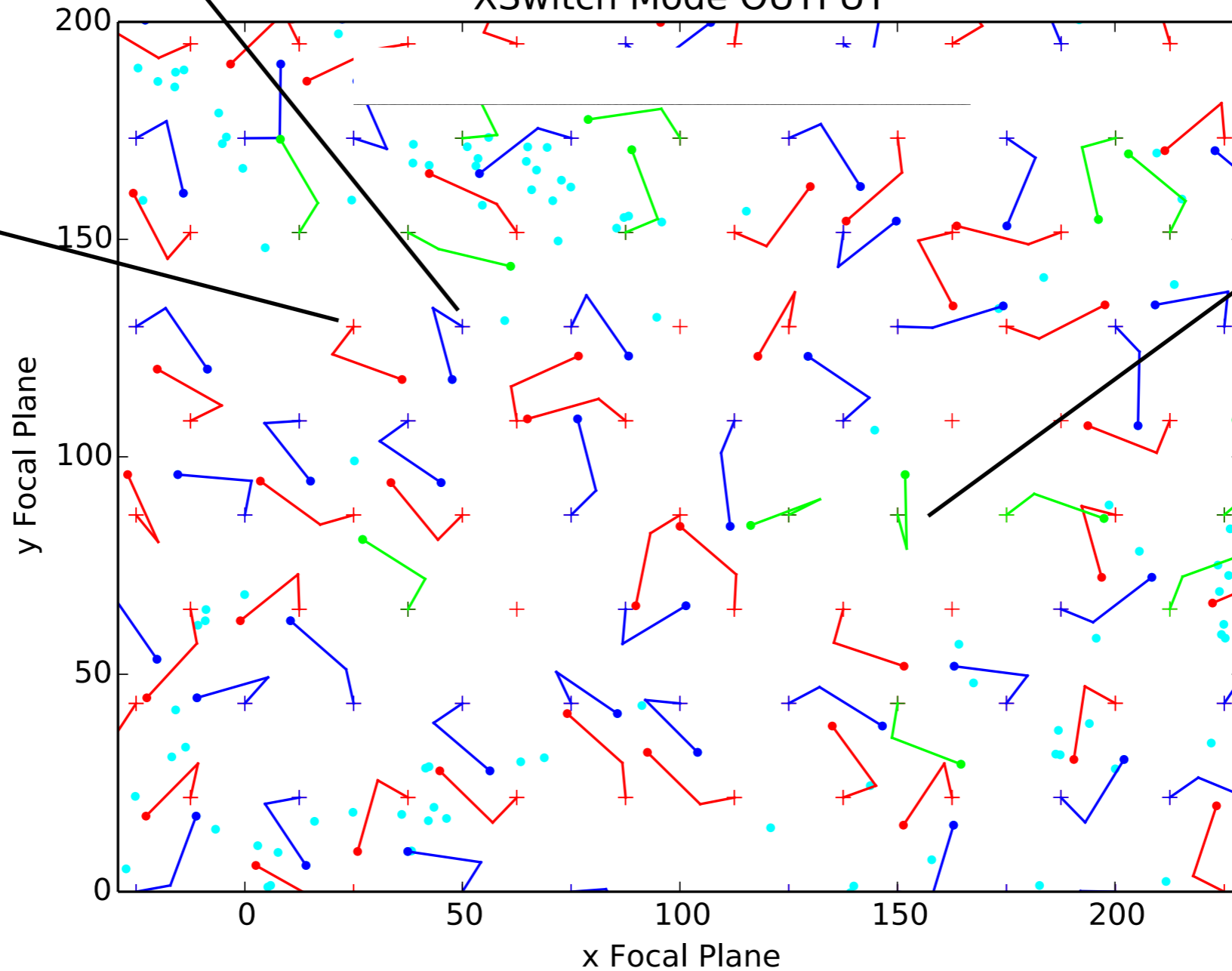
- TAR w/o SKY
- TAR not allocated
- SKY
- TAR with SKY



Fibre allocated
to a target

Fibre allocated
to a sky position

XSwitch Mode OUTPUT



Fibre allocated
to a target
(w/o sky)

nodding direction

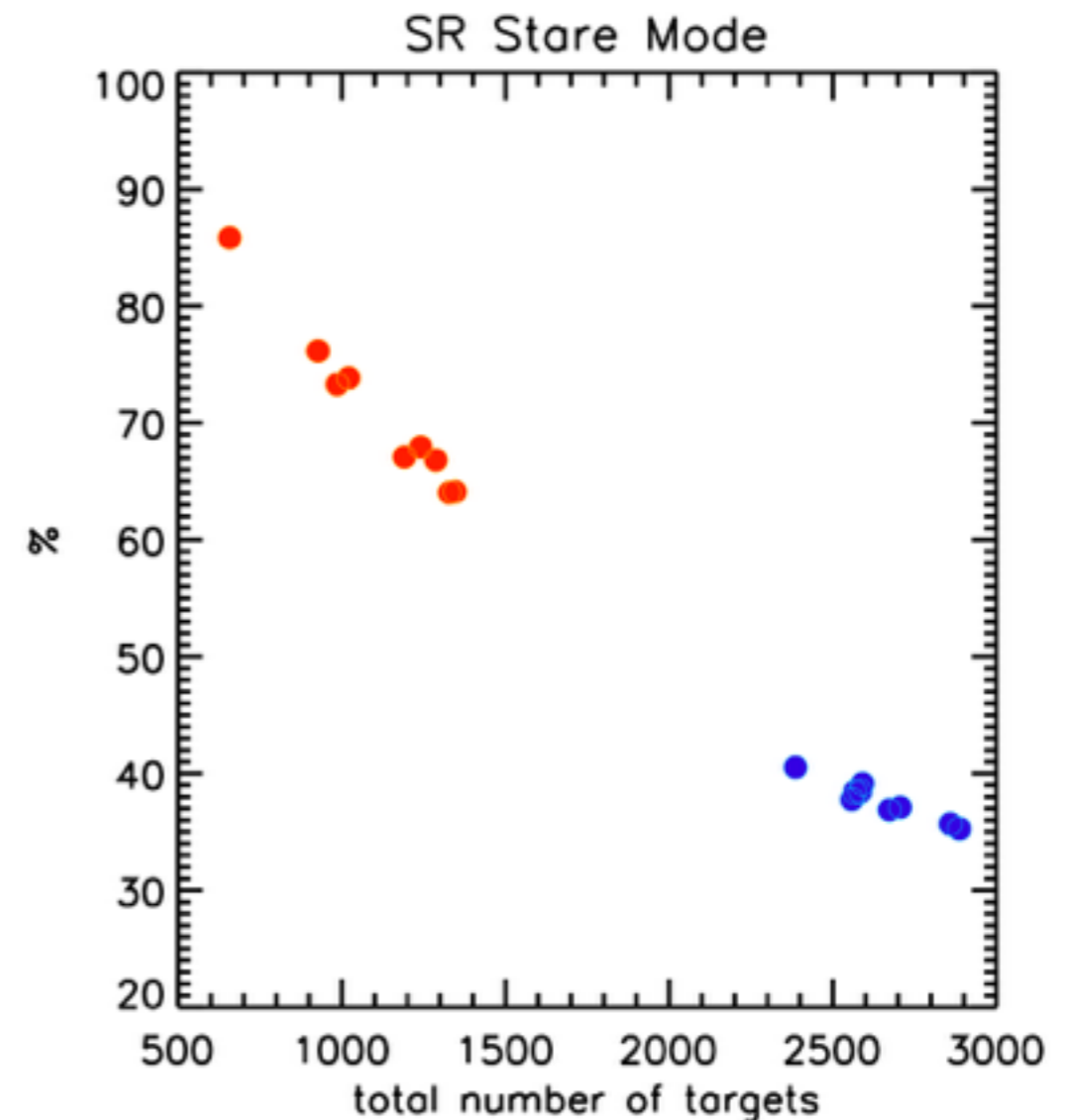
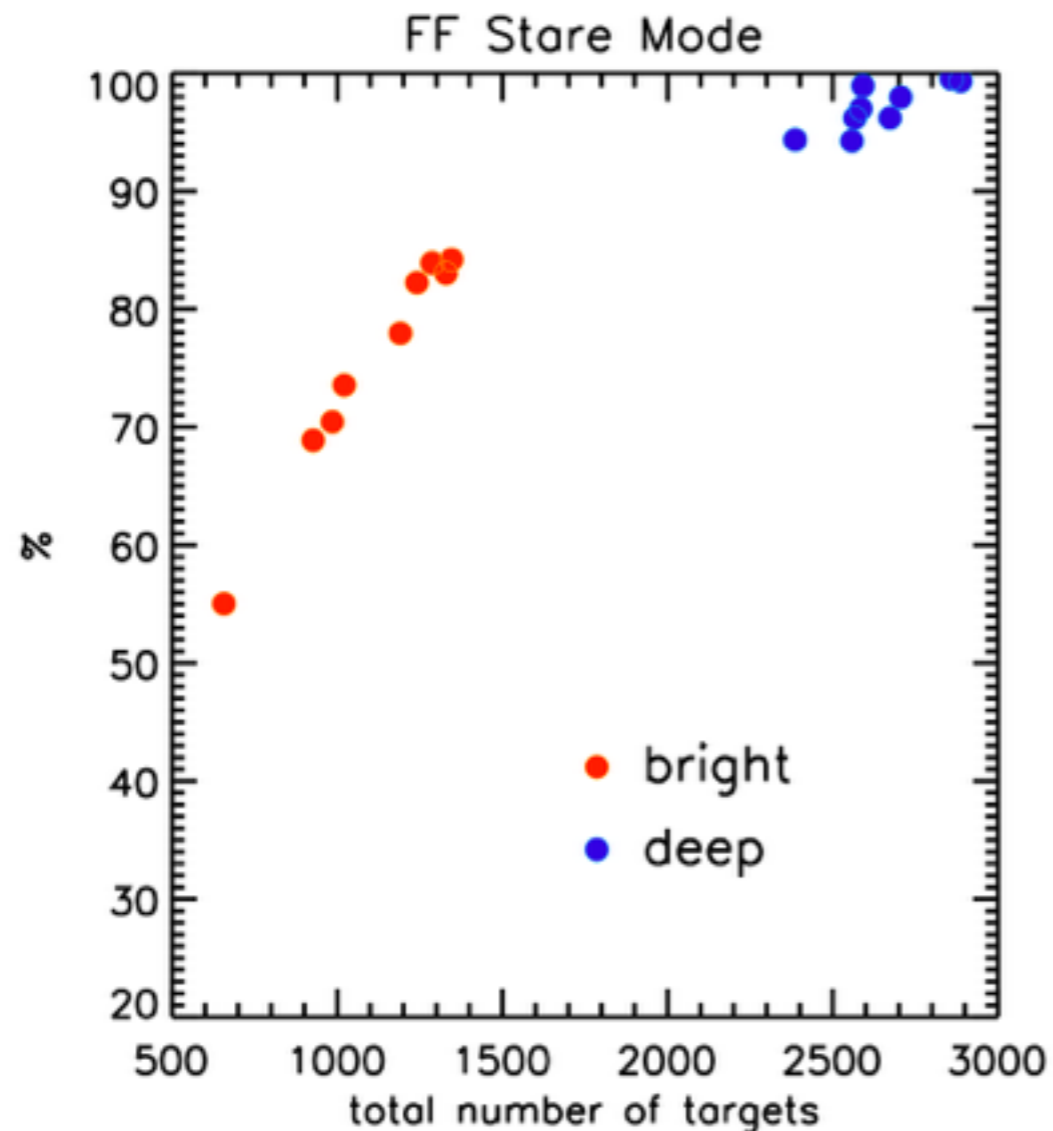
Statistics

Filling Factor (FF) :

number of fibres allocated / total number of fibres

Sampling Rate (SR) :

number of objects observed / total number of objects in the FoV



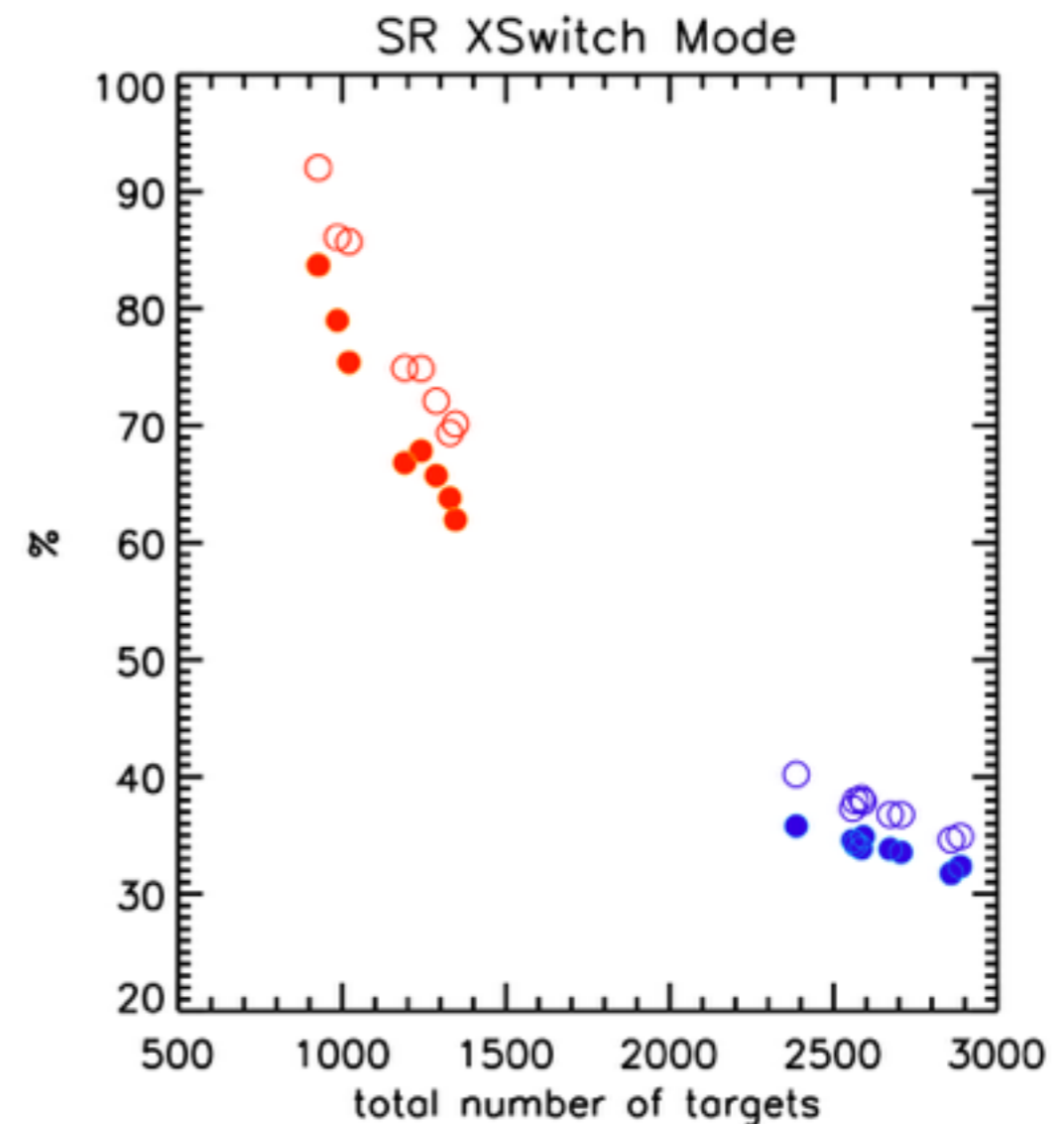
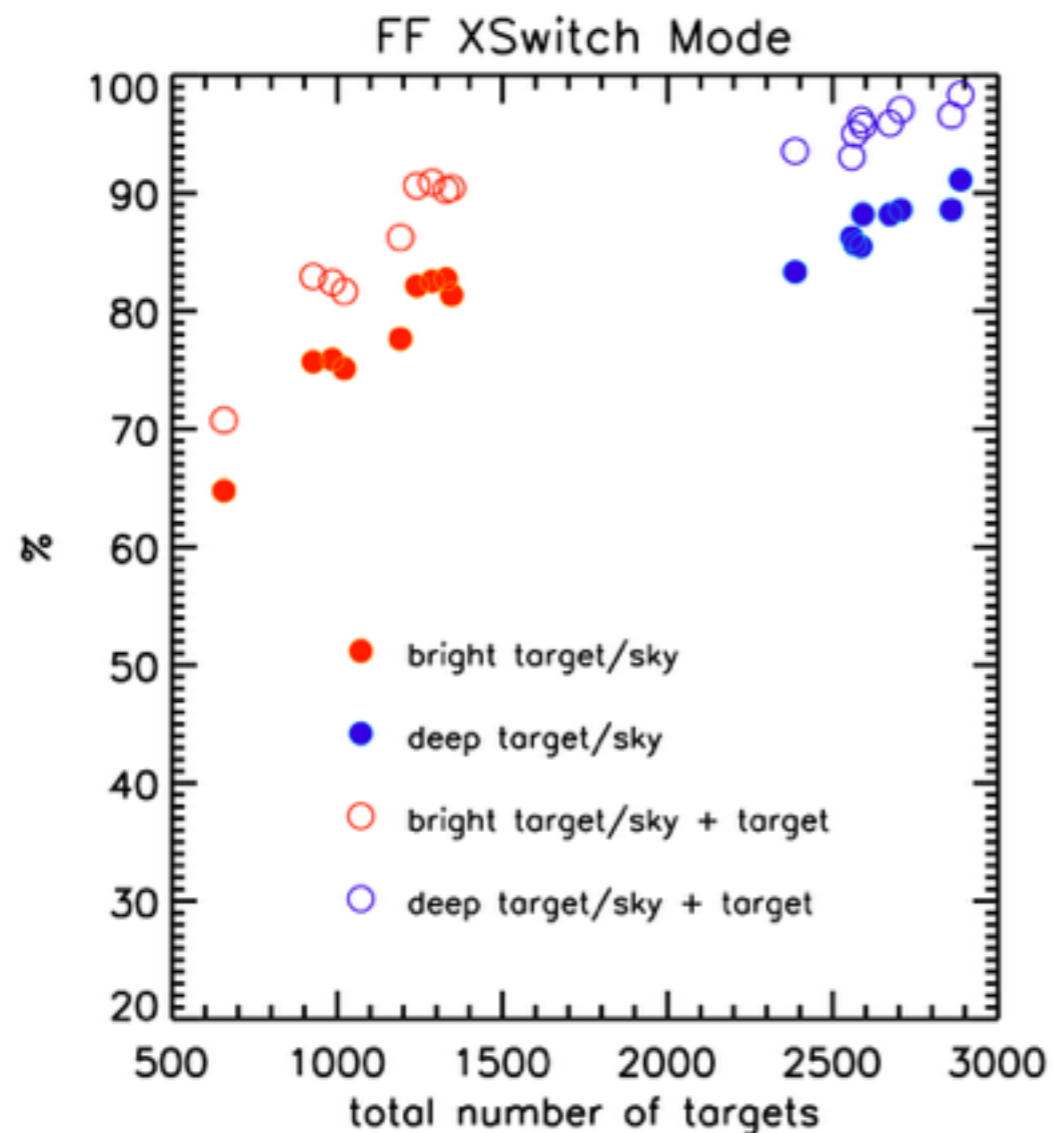
Statistics

Filling Factor (FF) :

number of fibres allocated / total number of fibres

Sampling Rate (SR) :

number of objects observed / total number of objects in the FoV



end